

# Galaxy populations from Deep *ISO* Surveys\*

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(11 August 2000)

**Abstract:** I discuss some of the main extra-galactic field surveys which have been undertaken by the Infrared Space Observatory (*ISO*). I review the findings from the source counts analyses and then examine some of the more recent detailed investigations into the explicit nature of the populations that make up these source counts.

## 1. *ISO* Surveys

Depending on how you count them, *ISO* (Kessler et al. 1996) undertook around fifteen independent field surveys to explore the extra-galactic populations. These surveys primarily used the *ISO*-PHOT (Lemke et al. (1996)) and *ISO*-CAM (Cesarsky et al. (1996)) instruments. As shown in Table 1 and illustrated in Figure 1 these surveys explored the full extent of *ISO*'s parameter space in terms of wavelength, area and depth.

This is an attempt to collect together some of the findings of these surveys. I will group the surveys by wavelength, since the populations explored at different wavelengths are expected to be somewhat different.

### 1.1 Mid IR surveys 6.7 - 15 $\mu$ m

Surveys in this wavelength regime tend to pick up emission from hotter sources such as AGN or luminous starbursts, they also pick up "cirrus" emission from normal galaxies. While these bands are located far from the peak of the typical starbursts spectral energy distribution it has been noted that the PAH features which do fall in these bands can be used as a good measure of star-formation activity. Surveys in these bands are the HDF surveys in the North and South; CFRS; ELAIS; the ISOCAM guaranteed time Extra Galactic Surveys, the Comet fields, the Japanese Lockman and Selected Area surveys and the CAM parallel mode survey.

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\* To appear in "Mid- and Far-infrared Astronomy and Future Space Missions", ISAS Report Special Edition (SP No. 14), ed. T. Matsumoto & H. Shibai

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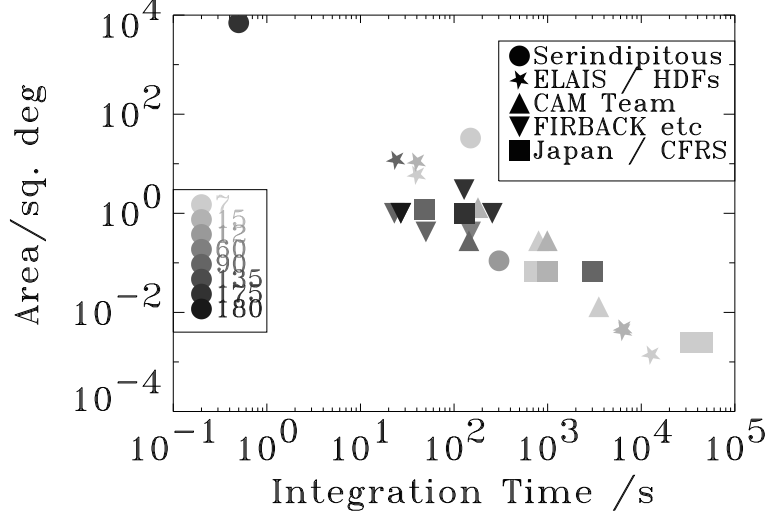


Fig. 1: Comparison of *ISO* surveys: area vs. depth. All the surveys in Table 1 are plotted. (Adapted from Oliver, et al. (2000b))

## 1.2 Far IR surveys 50-100 $\mu\text{m}$

These wavelengths are located at the peak of the rest-frame emission from obscured star formation and thus provide the most direct estimate of the obscured bolometric output in star-forming galaxies. Surveys in this band are that of SA57; the Japanese surveys; ELAIS; CIBR and CFRS.

## 1.3 “Cool” FIR surveys 120-200 $\mu\text{m}$

I use this term to denote bands beyond the normal peak in a star-burst emission. Surveys at these bands tend to pick up galaxies which are “cool” either because the energy distributions have been “cooled” simply by virtue of their high redshift, or because they are intrinsically cool but local. Surveys in this band are FIRBack; ELAIS, Japanese surveys; the PHOT slew surveys; and CIBR.

# 2. Counts

## 2.1 Mid Infrared

The ELAIS 15 $\mu\text{m}$  counts Serjeant et al. (2000), illustrated in Figure 2, were shown to agree with many strongly evolving models such as those of Franceschini et al. (1994), Pearson & Rowan-Robinson (1999) Xu et al. (1998) which had been based on *IRAS* counts, while being inconsistent with at least one non-evolving model that of Xu et al. (1998). Deeper counts such as those of the *ISO*-HDF Oliver et al. (1997) amplify this picture of strongly evolving populations and the summary presented by Elbaz et al. (1999) (also shown in Figure 2) suggests that the deepest cluster-lens

Table 1: Field Surveys with *ISO*, ordered roughly in decreasing area. References: 1 – Bogun et al. (1996), 2 – Siebenmorgen et al. (1996), 3 – Oliver, et al. (2000b) 4 – Elbaz et al. (1999), 5 – Dole et al. (1999), 6 – Juvela, et al. (2000), 7 – Linden-Vørnle (2000), 8 – Elbaz et al. (1999) 9 – Clements et al. (1996), 10 – Flores et al. (1999a), Flores et al. (1999b), 11 – Elbaz et al. (1999), 12 – Oliver et al. (2000, in preparation), 13 – Taniguchi et al. (1997), 14 – Kawara et al. (1998), 15 – Serjeant et al. (1997)

Survey Name	[e.g. ref]	Wavelength / $\mu\text{m}$	Integration /s	Area /sqdeg
PHT Serendipity Survey	1	175	0.5	7000
CAM Parallel Mode	2	6.7	150	33
ELAIS	3	6.7,15,90,175	40, 40, 24, 128	6, 11, 12,1
CAM Shallow	4	15	180	1.3
FIRBACK	5	175	256, 128	1, 3
CIBR	6	90, 135,180	23, 27, 27	1, 1, 1
SA 57	7	60, 90	150, 50	0.42,0.42
CAM Deep	8	6.7, 15, 90	800, 990, 144	0.28, 0.28, 0.28
Comet fields	9	12	302	0.11
CFRS	10	6.7,15,60,90	720, 1000, 3000,3000	0.067,0.067,0.067,0.067
CAM Ultra-Deep	11	6.7	3520	0.013
ISOHDF South	12	6.7, 15	> 6400, > 6400	4.7e-3, 4.7e-3
Deep Lockman	13,14	6.7, 90, 175	44640, 48, 128	2.5e-3, 1.2 , 1
ISOHDF North	15	6.7, 15	12800, 6400	1.4e-3, 4.2e-3

surveys have seen a turn over in the counts and thus the beginning of the end of these evolving populations.

## 2.2 FIR Counts

The FIR source counts also show good evidence for evolution. E.g. the ELAIS 90 $\mu\text{m}$  counts of Efstathiou et al. (2000) confirm the strong evolution expected from extrapolation from *IRAS* studies, see Figure 3. FIR counts from other *ISO* surveys are given in Juvela, et al. (2000) and Linden-Vørnle (2000).

## 2.3 “cool” FIR counts

The 175 $\mu\text{m}$  counts have attracted a lot of attention since they are more difficult to explain e.g. Dole et al. (1999) and references therein. The counts exceed models extrapolated from *IRAS*. To illustrate this Figure 4 shows the Dole et al. (2000) attempt to explain the “cool” counts. This model allowed a strong evolution in the ultra luminous *IRAS* galaxies and was constrained by both the 175 $\mu\text{m}$  counts and the FIR background. It should be stressed that such an extreme model is not a unique fit to the data, but illustrates the lengths to which one needs to go to explain these data.

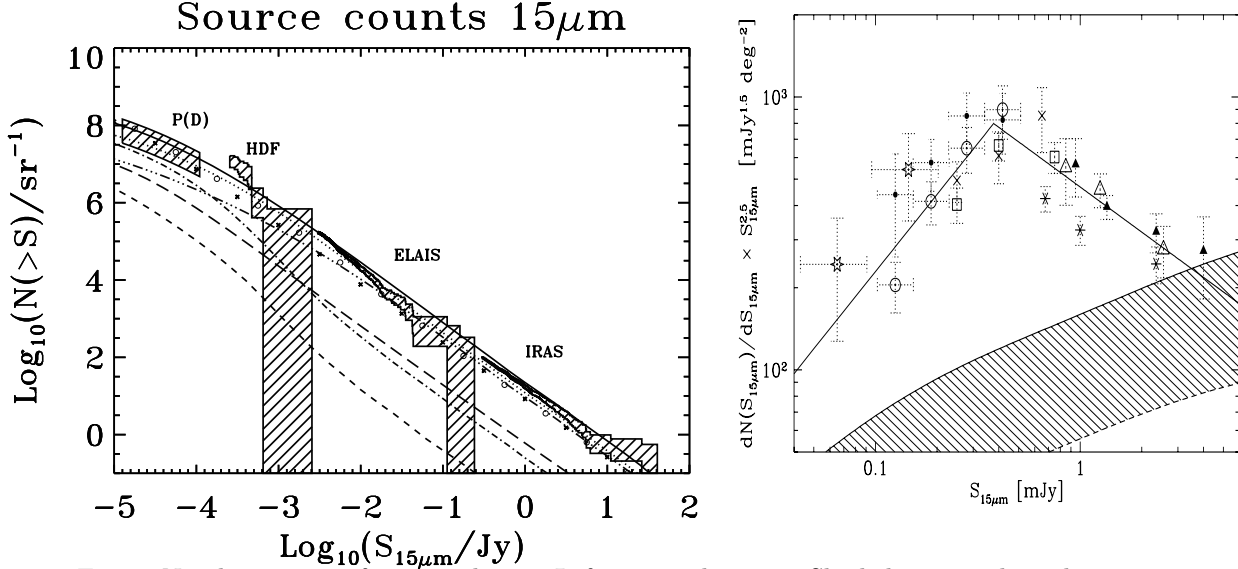


Fig. 2: Number counts of  $15\mu\text{m}$  galaxies. Left: integral counts. Shaded regions show the ranges spanned by  $\pm 1\sigma$  uncertainties. Also shown are the source counts and P(D) analysis from the Hubble Deep Field Oliver et al. (1997), Aussel et al. (1999). The *IRAS* counts are estimated from the  $12\mu\text{m}$  counts. The Franceschini et al. (1994) source count model is over-plotted. The model spiral contribution is shown as a dotted line, ellipticals as a dashed line, S0 as a dash-dot line, star-bursts as a dash-dot-dot-dot line and AGN as a long dashed line. The total population model is shown as a full line. Also plotted are the Guiderdoni et al. (1998) models A and E, as small filled crosses and small open circles respectively. Figure taken from Serjeant et al. (2000) Right: differential, Euclidean normalised counts. Data points: A2390 (open stars), *ISO*-HDF North (open circles), *ISO*-HDF-South (filled circles), Marano FIRBack (MFB) Ultra-Deep (open squares), Marano Ultra-Deep (crosses), MFB Deep (stars), Lockman Deep (open triangles), Lockman Shallow (filled triangles). The hatched area shows the range of expectations assuming no evolution. Figure Taken from Elbaz et al. (1999)

### 3. Beyond Counts

The counts are challenging, they show strong evolution, and may suggest new populations. However, in isolation they do not tell us what the nature of the sources is. The current effort is now strongly focused on more detailed investigations of these populations. I pick three particular results for discussion.

The first is from the ELAIS survey and is simply a summary of spectral classifications from the campaigns using 2dF and the ESO telescopes, Gruppioni et al. (in prep.) and La Franca et al. (in prep.), presented by Oliver, et al. (2000a) and shown in Table 2. The completeness and homogeneity of this list is not taken into account and so this table only gives a very crude picture. With that caveat in mind we do see that one third of these, relatively bright, mainly  $15\mu\text{m}$  sources have AGN of one sort or another. A second important result arises from the VLT spectroscopy of  $15\mu\text{m}$  *ISO*-HDF-South sources undertaken by Rigopoulou et al. 2000 and illustrated in Figure 5. This indicates that many of the fainter  $15\mu\text{m}$  sources, making up the peak in the differential counts are star-forming galaxies, hinting that the population mix may be changing as we move to fainter fluxes.

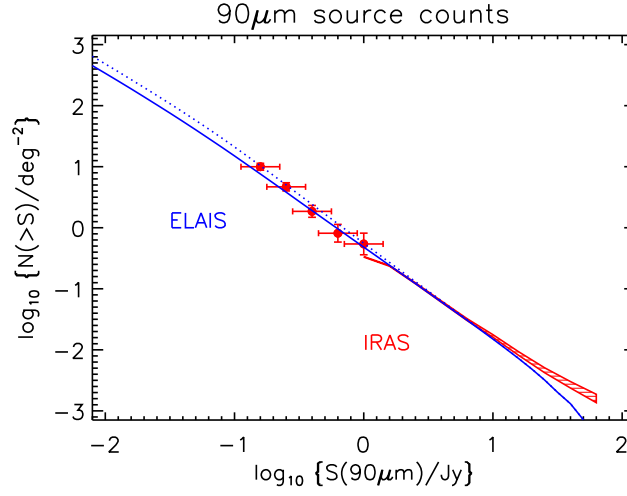


Fig. 3: ELAIS and *IRAS*  $90\mu\text{m}$  source counts. The solid line and dotted lines are the counts predicted by the models A and E (respectively) of Guiderdoni et al. (1998) Taken from Efstathiou et al. (2000)

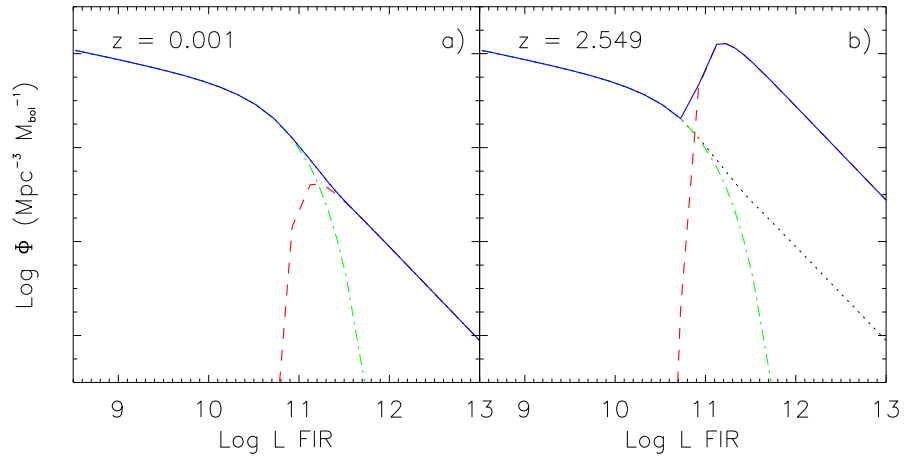


Fig. 4: Two snapshots of a model fitting the  $175\mu\text{m}$  counts and the FIR background, taken from Dole et al. (2000). *a*: Luminosity Function at  $z = 0$  (solid line); normal galaxy (dot-dash); ULIRG (dash-dash). *b*: Luminosity Function at  $z = 2.5$  (solid line); normal galaxy (dot-dash) ULIRG (dash-dash) and local LF (dots).

Table 2: Provisional spectral classifications of *ISO* and 21cm selected sources in the Southern ELAIS field S1. Spectra come from a one hour 2dF exposure and a number of nights on the ESO 3.6m and NTT telescopes (Gruppioni et al., in prep. and La Franca et al., in prep.). Table taken from Oliver, et al. (2000a)

Class	21cm		<i>ISO</i>	
	2dF	ESO	2dF	ESO
Absorption	48		15	6
Star-burst	9	3	22	52
H $\alpha$	5	2	26	
OII	10		4	
OIII	1			
AGN/QSO	6	2	20	19
AGN/Sy1	2	1	3	
AGN/Sy2	8	3	8	8
AGN/BLLac			2	
Stars	1		8	3
Too Faint	60	3	41	

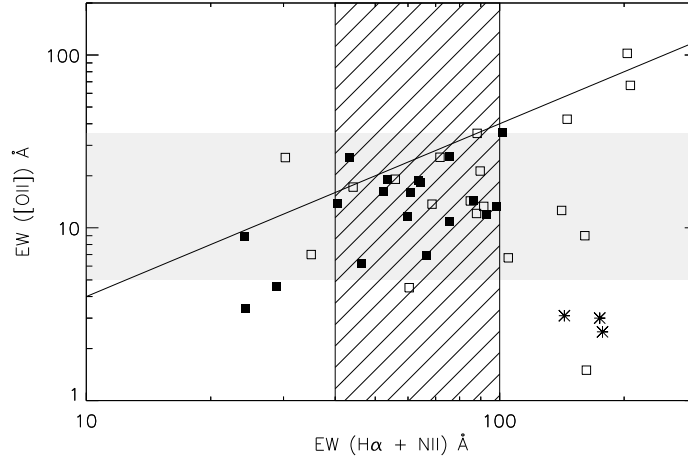


Fig. 5:  $EW(OII)$  vs.  $EW(H\alpha+NII)$ , filled squares e(a) galaxies, open squares non-e(a) galaxies, stars Seyferts). The black line corresponds to  $EW(OII)=0.4 \cdot EW(H\alpha+NII)$  found for nearby field galaxies by Kennicutt (1992). The hatched vertical region shows the location of VLT measurements of the *ISO*-HDF-S galaxies, the shaded horizontal region the range from the *ISO*-CFRS sources. The intersection of the two bars represents the inferred location of the VLT *ISO*-HDF-S galaxies, indistinguishable from the dusty, luminous, e(a) galaxies. Figure taken from Rigopoulou et al. 2000

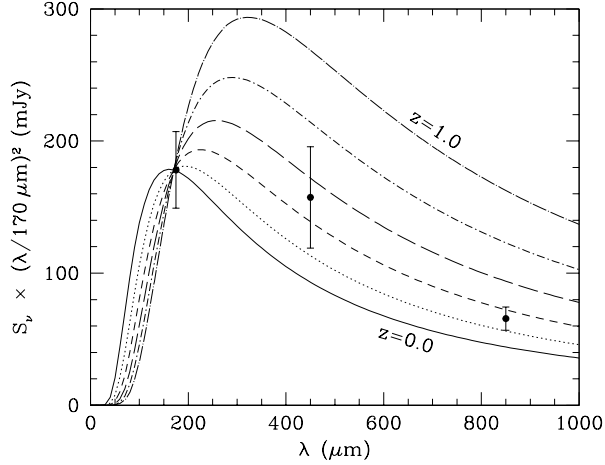


Fig. 6: Average spectral energy distribution for a FIRBACK sub-sample, from Scott et al. (2000). The  $170\mu\text{m}$  point is the average of 10 FIRBACK sources, with error-bar being the standard error from the scatter among them. The  $450\mu\text{m}$  and  $850\mu\text{m}$  points are from SCUBA observations. For the sake of visual clarity  $y$ -axis has been multiplied by  $x^2$ . The curves show emission from modified blackbodies, normalised to the  $170\mu\text{m}$  flux density, with  $T_d = 40\text{ K}$  and  $\beta = 1.5$ , and for  $z = 0.0$  (solid line) up to  $z = 1.0$  in steps of  $0.2$ . Note that the shapes of these curves are degenerate in the combination  $(1 + z)/T_d$ .

The third result comes from studies of the  $175\mu\text{m}$  sources. These have been particularly hard to identify spectroscopically but the sub-mm photometric study of Scott et al. (2000) is instructive. Figure 6 shows the average fluxes of 10 sources compared with some simply SEDs. From this it appears that the best fitting models would have the sources at relatively modest redshifts.

## 4. CONCLUSIONS

### 4.1 What *ISO* source counts tell us

- Strong significant evolution seen at all wavelengths observable to *ISO*.
- Differential  $15\mu\text{m}$  counts appear to steepen and then decline below about  $1\text{ mJy}$
- Source counts particularly steep at  $175\mu\text{m}$
- *ISO* reached the confusion limit at all wavelengths, easily at the longest wavelengths
- 10% of the FIR background was resolved at long wavelengths
- As much as 30% of the FIR background may come from the sources detected at mid IR wavelengths.
- Many models exist which might explain counts, though  $175\mu\text{m}$  counts are harder to explain than other bands

## 4.2 What *ISO* survey follow-up tells us

- Fainter mid IR sources appear to be predominantly star-forming, rather than AGN
- Brighter mid IR sources may contain have a higher fraction of AGN
- Best current guess is that  $175\mu\text{m}$  sources are  $z < 1$  although based on limited evidence

Forthcoming followup projects, notably X-ray surveys in *ISO* survey fields, should help enormously to identify what populations make up these strongly evolving *ISO* sources.

## ACKNOWLEDGMENT

This paper is based on observations with *ISO*, an ESA project, with instruments funded by ESA Member States (especially the PI countries: France, Germany, the Netherlands and the United Kingdom) and with participation of ISAS and NASA. This work was in part supported by PPARC grant no. GR/K98728 and EC Network is FMRX-CT96-0068.

I would like to thank the local organising committee for all the efforts and for making this a very enjoyable meeting.

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